



Periodic Radio and H α emission from the L dwarf binary 2MASSW J0746425+200032: exploring the magnetic field topology and radius of an L dwarf

Citation

Berger, E., R. E. Rutledge, N. Phan-Bao, G. Basri, M. S. Giampapa, J. E. Gizis, J. Liebert, E. Martín, and T. A. Fleming. 2009. Periodic Radio and H α emission from the L dwarf binary 2MASSW J0746425+200032: exploring the magnetic field topology and radius of an L dwarf. *The Astrophysical Journal* 695, no. 1: 310–316. doi:10.1088/0004-637x/695/1/310.

Published Version

doi:10.1088/0004-637x/695/1/310

Permanent link

<http://nrs.harvard.edu/urn-3:HUL.InstRepos:30410818>

Terms of Use

This article was downloaded from Harvard University's DASH repository, and is made available under the terms and conditions applicable to Open Access Policy Articles, as set forth at <http://nrs.harvard.edu/urn-3:HUL.InstRepos:dash.current.terms-of-use#OAP>

Share Your Story

The Harvard community has made this article openly available.
Please share how this access benefits you. [Submit a story](#).

[Accessibility](#)

PERIODIC RADIO AND H α EMISSION FROM THE L DWARF BINARY 2MASSW J0746425+200032: EXPLORING THE MAGNETIC FIELD TOPOLOGY AND RADIUS OF AN L DWARF

E. BERGER¹, R. E. RUTLEDGE², N. PHAN-BAO⁹, G. BASRI⁴, M. S. GIAMPAPA⁵, J. E. GIZIS⁶, J. LIEBERT⁷, E. MARTÍN^{8,9}, AND T. A. FLEMING⁷,

Draft version August 29, 2008

ABSTRACT

We present an 8.5-hour simultaneous radio, X-ray, UV, and optical observation of the L dwarf binary 2MASSW J0746425+200032. We detect strong radio emission, dominated by short-duration periodic pulses at 4.86 GHz with $P = 124.32 \pm 0.11$ min. The stability of the pulse profiles and arrival times demonstrates that they are due to the rotational modulation of a $B \approx 1.7$ kG magnetic field. A quiescent non-variable component is also detected, likely due to emission from a uniform large-scale field. The H α emission exhibits identical periodicity, but unlike the radio pulses it varies sinusoidally and is offset by exactly 1/4 of a phase. The sinusoidal variations require chromospheric emission from a large-scale field structure, with the radio pulses likely emanating from the magnetic poles. While both light curves can be explained by a rotating mis-aligned magnetic field, the 1/4 phase lag rules out a symmetric dipole topology since it would result in a phase lag of 1/2 (poloidal field) or zero (toroidal field). We therefore conclude that either (i) the field is dominated by a quadrupole configuration, which can naturally explain the 1/4 phase lag; or (ii) the H α and/or radio emission regions are not trivially aligned with the field. Regardless of the field topology, we use the measured period along with the known rotation velocity ($v \sin i \approx 27$ km s⁻¹), and the binary orbital inclination ($i \approx 142^\circ$), to derive a radius for the primary star of 0.078 ± 0.010 R $_{\odot}$. This is the first measurement of the radius of an L dwarf, and along with a mass of 0.085 ± 0.010 M $_{\odot}$ it provides a constraint on the mass-radius relation below 0.1 M $_{\odot}$. We find that the radius is about 30% smaller than expected from theoretical models, even for an age of a few Gyr. The origin of this discrepancy is either a breakdown of the models at the bottom of the main sequence, or a significant mis-alignment between the rotational and orbital axes.

Subject headings: radio continuum:stars — stars:activity — stars:low-mass, brown dwarfs — stars:magnetic fields

1. INTRODUCTION

Radio observations conducted over the past several years have uncovered a substantial fraction of magnetically active low mass stars and brown dwarfs. Both quiescent and flaring emission are present, with luminosities that remain unchanged down to at least spectral type \sim L3, in contrast to the declining activity seen in X-rays and H α (Berger et al. 2001; Berger 2002; Berger et al. 2005; Burgasser & Putman 2005; Berger 2006; Osten et al. 2006; Hallinan et al. 2006; Antonova et al. 2007; Phan-Bao et al. 2007; Hallinan et al. 2007; Audard et al. 2007; Antonova et al. 2008; Berger et al. 2008a,b; Hallinan et al. 2008). Depending on the nature of the radio emission mechanism, the inferred magnetic field strengths are $\sim 0.1 - 3$ kG with order unity filling factors (Berger 2006; Hallinan et al. 2008). Long term monitoring of several objects in the spectral type range M8–L3.5 has further shown that the fields are generally stable for at least

several years (Berger et al. 2005; Berger 2006; Berger et al. 2008a; but see also Antonova et al. 2007), providing an important constraint on the lifetime of magnetic dynamos in fully convective objects. The strength, scale, and stability of the inferred fields is in good agreement with recent results from phase-resolved spectropolarimetry (spectral type \sim M4; Donati et al. 2006; Morin et al. 2008) and Zeeman broadening of FeH lines (spectral type $<$ M9; Reiners & Basri 2006, 2007), as well as with the most recent numerical dynamo simulations (e.g., Browning 2008).

Equally important, three ultracool dwarfs to date have been observed to produce periodic radio emission, with periods of 184 min (2MASS J00361617+1821104; Berger et al. 2005), 118 min (TVLM 513-46546⁹; Hallinan et al. 2007), and 170 min (LSR 1835+3259; Hallinan et al. 2008). In the case of 2M 0036+18 the periodic emission is sinusoidal, while in the latter two objects it is in the form of short duration pulses (duty cycle of a few percent). In all three cases the observed periods are in good agreement with the known rotation velocities ($v \sin i$), indicating that the radio periodicity traces the stellar rotation.

As a result, in addition to allowing a measurement of the magnetic field properties, radio observations provide a unique opportunity to measure the radii of late-M and L dwarfs through the combination of rotation period and velocity. However, since only $v \sin i$ values are known for these objects, there is an inherent degeneracy between the radius and the inclination of the rotation axis, i . Using a typical range of

¹ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138

³ Institute of Astronomy and Astrophysics, Academia Sinica, P.O. Box 23-141, Taipei 10617, Taiwan, R.O.C.

⁴ Astronomy Department, University of California, Berkeley, CA 94720

⁵ National Solar Observatory, National Optical Astronomy Observatories, Tucson, AZ 85726

⁶ Department of Physics and Astronomy, University of Delaware, Newark, DE 19716

⁷ Department of Astronomy and Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721

⁸ Instituto de Astrofísica de Canarias, C/ Vía Láctea s/n, E-38200 La Laguna, Tenerife, Spain

⁹ University of Central Florida, Department of Physics, PO Box 162385, Orlando, FL 32816

⁹ This object also exhibits periodic H α emission, with the same period as observed in the radio (Berger et al. 2008a).

$R \approx 0.09\text{--}0.11 R_{\odot}$ for late-M and L dwarfs, the allowed range of inclinations span a relatively wide range of $\sim 60\text{--}90^{\circ}$ (Berger et al. 2005; Hallinan et al. 2008). To break this degeneracy, and thus measure the radius directly, it is desirable to observe objects for which the inclination can be estimated. This is the case for binary systems with well determined orbital parameters if we make the reasonable assumption that the rotation and orbital axes are aligned (Hale 1994). Along with an estimate of the mass, we can thus place constraints on the mass-radius relation for ultracool dwarfs.

Here we present simultaneous radio, X-ray, optical, and UV observations of one such system, the L dwarf binary 2MASSW J0746425+200032 (hereafter, 2M 0746+20). These observations are part of a long-term project to study the field properties of ultracool dwarfs (Berger et al. 2005, 2008a,b). While no X-ray or UV emission are detected, the radio and $H\alpha$ emission exhibit clear periodicity, which along with the known orbital inclination and $v \sin i$ allow us to explore for the first time the radius and magnetic field topology of an L dwarf.

2. OBSERVATIONS

2.1. Target Selection and Properties

We targeted the L dwarf binary 2M 0746+20 (L0+L1.5) due to the availability of its orbital parameters and rotation velocity, as well as previous detections in $H\alpha$. The binary is located at a distance of only 12.2 pc (Dahn et al. 2002). The orbital inclination and total binary mass are $142 \pm 3^{\circ}$ and $0.146^{+0.016}_{-0.006} M_{\odot}$, respectively (Bouy et al. 2004). The rotation velocities quoted in the literature are $26 \pm 3 \text{ km s}^{-1}$ (Bailer-Jones 2004), 24 km s^{-1} (Reid et al. 2002), 31 km s^{-1} (Reiners & Basri 2008), and 28 km s^{-1} (C. Blake priv. comm.) We therefore adopt an average value of $v \sin i = 27 \pm 3 \text{ km s}^{-1}$. The bolometric luminosity of 2M 0746+20 is $L_{\text{bol}} = 10^{-3.64 \pm 0.06} L_{\odot}$ (Vrba et al. 2004). We note that there is still no agreement as to whether the secondary member of the binary is a low mass star or a brown dwarf (Bouy et al. 2004; Gizis & Reid 2006).

Previous $H\alpha$ detections revealed an equivalent width range of $\approx 1.2\text{--}2.4 \text{ \AA}$, or $\log(L_{H\alpha}/L_{\text{bol}}) \approx -5.3$ to -5.2 (Reid et al. 2000, 2002; Schmidt et al. 2007; Reiners & Basri 2008). No radio emission was previously detected at 8.46 GHz to a 3σ limit of $\lesssim 48 \mu\text{Jy}$ (Berger 2006), but subsequent to the observations presented here, Antonova et al. (2008) published a radio detection at 4.86 GHz with $F_{\nu} = 286 \pm 24 \mu\text{Jy}$, based on a 2-hour observation. They further detect the possible emergence of a flare in the last few minutes of their observation.

The simultaneous observations presented here were conducted on 2008 February 22 UT for a total of 8.4 hr in the radio (02:11–10:34 UT), 8.83 hr in the X-rays (02:20–11:10 UT), and 7.4 hr in the optical (05:43–13:09 UT). Observations with the *Swift* UV/Optical Telescope (UVOT) took place intermittently between 01:32 and 09:54 UT with a total on-source exposure time of 2.15 hr.

2.2. Radio

Very Large Array¹⁰ observations were obtained simultaneously at 4.86 and 8.46 GHz in the standard continuum mode with $2 \times 50 \text{ MHz}$ contiguous bands. Thirteen antennas were used at each frequency in the BnC array configuration. Scans

of 295 s on source were interleaved with 55 s scans on the phase calibrator J0738+177. The flux density scale was determined using the extragalactic source 3C 286 (J1331+305). Data reduction and analysis follow the procedures outlined in Berger et al. (2008a) and Berger et al. (2008b). We detect a source coincident with the position¹¹ of 2M 0746+20 at both frequencies.

2.3. Optical Spectroscopy

We used¹² the Gemini Multi-Object Spectrograph (GMOS; Hook et al. 2004) mounted on the Gemini-North 8-m telescope with the B600 grating at a central wavelength of 5250 \AA , and with a $1''$ slit. The individual 300-s exposures were reduced using the *gemini* package in IRAF (bias subtraction, flat-fielding, and sky subtraction), and the wavelength solution was determined from CuAr arc lamps. The spectra cover 3840–6680 \AA at a resolution of about 5 \AA . A total of 77 exposures were obtained, with an on-source efficiency of 94%. We detect $H\alpha$ emission in all the individual spectra.

2.4. X-Rays

Observations were performed with the Chandra/ACIS-S3 backside-illuminated chip for a total of 29.46 ks. The data were analyzed using CIAO version 3.4, and counts were extracted in a $2''$ radius circle centered on the position of 2M 0746+20. We find only 2 counts, with 1.5 counts expected from the background as determined from annuli centered on the source position. Thus, the resulting upper limit is about 7 counts (95% confidence level). Using an energy conversion factor of $1 \text{ cps} = 3.8 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ (appropriate for a 1 keV Raymond-Smith plasma model in the 0.2–2 keV range) we find $F_X < 9.0 \times 10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1}$, or $L_X/L_{\text{bol}} \lesssim 10^{-4.7}$.

2.5. Ultraviolet

Data were obtained with the *Swift*/UVOT in the UVW1 filter ($\lambda_{\text{eff}} \approx 2600 \text{ \AA}$), as a series of 6 images with exposure times ranging from 460 to 1625 s. No source is detected at the position of 2M 0746+20 in any of the individual exposures or in the combined 2.15 hr image. Photometry of the combined image results in a limit of $F_{\lambda} < 2.8 \times 10^{-18} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ in a $2''$ aperture. The ratio of UV to bolometric luminosity is $\lambda L_{\lambda}/L_{\text{bol}} \lesssim 10^{-3.8}$.

3. MULTI-WAVELENGTH PERIODIC EMISSION

3.1. Radio Emission

The radio emission from 2M 0746+20 has an average flux density of $304 \pm 15 \mu\text{Jy}$ at 4.86 GHz, and $154 \pm 14 \mu\text{Jy}$ at 8.46 GHz. This is the first detection of the object at 8.46 GHz, with an increase by at least a factor of three compared to previous limits from June 2002 (Berger 2006). The average fraction of circular polarization is $\lesssim 15\%$ (3σ) and $35 \pm 10\%$, respectively. The resulting average luminosities are $L_{\nu}(4.86) = (5.4 \pm 0.3) \times 10^{13}$ and $L_{\nu}(8.46) = (2.8 \pm 0.3) \times 10^{13} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$, and the ratio relative to the bolometric luminosity is $\nu L_{\nu}/L_{\text{bol}} \approx 10^{-6.55}$. These values are similar to those of previously-detected late-M and L dwarfs (Berger 2006).

The radio light curves are shown in Figure 1. The most striking aspect of the radio emission is a set of bright (10–15

¹⁰ The VLA is operated by the National Radio Astronomy Observatory, a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

¹¹ We take into account the known proper motion of 370 mas yr^{-1} at a position angle of 264° (Schmidt et al. 2007).

¹² Observations were obtained as part of program GN-2008A-Q-11.

mJy), short duration (1.2 min), circularly polarized ($\sim 100\%$), and *periodic* pulses at 4.86 GHz. The period measured from the well-defined peaks¹³ of the four detected pulses is 124.32 ± 0.11 min. No corresponding emission is detected at 8.46 GHz, indicating that the fractional bandwidth is $\delta\nu/\nu \lesssim 0.7$. The pulses are detected at both intermediate frequencies of the 4.86 GHz band (4885 MHz and 4835 MHz), with no clear time delay implying that $\delta\nu/\nu \gtrsim 0.02$.

The detailed temporal profiles of the pulses in total intensity and circular polarization are shown in Figure 2. The profiles are similar in all four cases, with a rise time of about 20 s, followed by a decline timescale of 40 s. The high degree of stability of the period and emission properties demonstrate that the pulsing activity is the result of stellar rotation, rather than genuine episodic flares. In this context, the duty cycle of only 0.8% implies that the radio emitting region has an azimuthal scale of only $\sim 0.1 R_*$ (assuming a height above the surface of about $0.5 - 1 R_*$; Hallinan et al. 2008). Furthermore, since the measured period roughly agrees with the rotation velocity of 2M 0746+20 (§2), we conclude that only one pulse is observed per rotation. Along with the uniform sense of circular polarization, this indicates that we observe only one of the magnetic poles.

The properties of the periodic pulses are similar to those observed in TVLM 513-46546 and LSR 1835+32 (Hallinan et al. 2007, 2008), and point to a coherent emission process, most likely the electron cyclotron maser (e.g., Treumann 2006). In this framework, the radiation is produced primarily at the electron cyclotron frequency, $\nu_c \approx 2.8 \times 10^6 B$ Hz, indicating that the magnetic field strength of 2M 0746+20 is $B \approx 1.7$ kG. We can also place a limit on the electron density using the constraint that the plasma frequency, $\nu_p \approx 9 \times 10^3 n_e^{1/2}$ Hz, has to be lower than ν_c . We thus find, $n_e \lesssim 3 \times 10^{11} \text{ cm}^{-3}$.

A single short duration burst is also detected at 8.46 GHz (03:49:36 UT; Figure 1), with a peak flux of about 6 mJy, an identical duration to the 4.86 GHz pulses, and $\sim 100\%$ right circular polarization. No corresponding emission is detected at 4.86 GHz. In the context of electron cyclotron maser emission, this burst requires a stronger magnetic field component of $B \approx 3$ kG. The burst is delayed by 72 min (or $0.58 \times P$) relative to the preceding 4.86 GHz pulse, indicating that its emission region is not trivially related to that of the periodic pulses. This, and the lack of periodicity, suggest that the 8.46 GHz burst is either the result of transitory field dissipation (a genuine flare), or that the physical conditions in its emission region (field strength and/or density) vary rapidly with time so as to suppress the emerging 8.46 GHz radiation during subsequent rotations. In either case, it is clear that the magnetic field strength ranges by at least a factor of two over the stellar surface and/or corona.

Eliminating the contribution of the short duration pulses, we find baseline quiescent emission of $224 \pm 15 \mu\text{Jy}$ (4.86 GHz) and $149 \pm 15 \mu\text{Jy}$ (8.46 GHz) with no clear sign of variability (Figure 1). The quiescent emission is distinct from the narrow-band pulses since it is detected at both 4.86 and 8.46 GHz and it does not vary with the stellar rotation. The wide frequency range ($\delta\nu/\nu \sim 1$) and inferred spectral index of -0.7 ± 0.3 are typical of gyrosynchrotron emission from a power law distribution of electrons with an energy index of $p \approx 2.4$ (Güdel 2002). The lack of variability during 8.4

hours (or four rotations) suggests that the quiescent emission arises from particle acceleration in a uniform large-scale field component with order unity covering fraction.

3.2. H α Emission

The H α light curve is shown in Figure 3. The equivalent width ranges from 2.4 to 3.1 Å, with a mean value of about 2.7 Å. More importantly, the light curve is clearly periodic, with $P = 126 \pm 10$ min in excellent agreement with the radio periodicity. This indicates that the radio and H α emission arise from the same binary member. However, despite the identical periods, two clear differences are present between the H α and radio light curves, which provide additional constraints on the geometry of the magnetic field.

First, the H α light curve is sinusoidal as opposed to the $\sim 1\%$ duty cycle of the radio pulses. This indicates that the H α emission is produced by chromospheric plasma that covers a substantial fraction (though less than 100%) of the stellar surface and rotates relative to our line of sight. A small enhancement is ruled out since it would result in a sharp rise and decline through ingress and egress, respectively, and a flat-topped peak. In addition, the non-zero minimum equivalent width (≈ 2.4 Å) indicates that the sinusoidal variations are most likely due to a combination of rotation and field orientation effects, such that at any given time some fraction of the H α -emitting chromosphere is visible, with the maximum projected solid angle corresponding to the light curve peaks (see §4).

Second, while the radio and H α periods are identical, the phase of the two light curves differs (Figure 3). In particular, we find that the H α peaks lead the radio pulses by about 31 min, corresponding to 1/4 of the period (or, equivalently 1/4 of a rotation). This lag indicates that the narrow beam which gives rise to the radio pulses is offset by about 90° relative to the central axis of the chromospheric geometry. As we show in the following section, such a 90° offset cannot be accommodated in a simple dipole magnetic field model, unless there is a significant breaking of the symmetry between the field and emission regions.

We have previously observed similar periodic H α emission from the M8.5 dwarf TVLM 513-46546, with $P \approx 2$ hr well-matched to the rotation velocity of $v \sin i \approx 60 \text{ km s}^{-1}$ (Berger et al. 2008a). This object has also been shown to produce radio bursts on a separate occasion (Hallinan et al. 2007). In the case of LSR 1835+32, on the other hand, despite the presence of periodic radio bursts (Hallinan et al. 2008) no periodic H α emission was evident in a separate 5.4 hr observation (Berger et al. 2008b). It is thus possible that the magnetic field topology and stability timescale, as well as the viewing and magnetic axis orientation, play a role in determining the correlation (or lack thereof) between the periodic radio and H α signals.

4. THE MAGNETIC FIELD GEOMETRY

Taking into account the properties of the radio and H α emission, we now investigate the magnetic field topology. This topology has to satisfy the following requirements: (i) the H α emission is due to a large scale feature; (ii) the non-zero minimum of the H α light curve requires some fraction of the active chromosphere to be visible at all times; (iii) the short duration radio pulses arise from a narrow region; and (iv) the radio pulses lag the H α peaks by 1/4 of the period.

The first scenario we explore is a poloidal field with a simple dipole geometry. Such fields have been inferred for

¹³ The UT times of the four peaks are 02:37:37.6, 04:41:57.5, 06:46:07.3, and 08:50:37.7.

convective mid-M dwarfs, such as V374 Peg, from phase-resolved spectropolarimetry (Donati et al. 2006; Morin et al. 2008). This field geometry, and our line of sight orientation, are shown in Figure 4. We assume that the inclination of the rotation axis relative to our line of sight is identical to the orbital axis, $i = 142^\circ$. In order to explain the radio pulses with a small duty cycle the magnetic axis has to be tilted relative to the rotation axis so that the beam emerging from the radial field at the poles sweeps into our line of sight once per rotation (a “pulsar configuration”). This configuration also explains the $H\alpha$ modulation since the solid angle subtended by the poloidal field that is projected along our line of sight varies sinusoidally as the object rotates. Specifically, as the field is tilted towards our line of sight we observe maximum $H\alpha$ emission (see panels A and B of Figure 4 for the geometry corresponding to the maximum and minimum $H\alpha$ emission, respectively).

Unfortunately, while this simple model explains both the radio pulsations and the $H\alpha$ light curve, it predicts a $1/2$ phase lag between the two light curves, rather than the observed $1/4$ phase lag. This is simply because the same phase of the rotation that orients the magnetic pole toward our line of sight (resulting in a radio burst), also projects the minimum solid angle of the dipole field (i.e., minimum $H\alpha$ emission). This effect is independent of our choice of orbital inclination (Figure 4).

Alternatively, a toroidal dipole configuration, with the polar caps producing the brightest $H\alpha$ emission, will also produce radio bursts and sinusoidal $H\alpha$ emission. However, in this configuration the two light curves would be exactly in phase since when the polar region is oriented toward our line of sight we will observe both an enhancement of the $H\alpha$ emission and a radio pulse.

Having rejected the simple dipole poloidal and toroidal configurations, we are left with two possibilities to explain the $1/4$ phase lag and the radio/ $H\alpha$ light curves. First, the magnetic field configuration is more complex and dominated by higher order multipoles. Or second, there is a breaking of the symmetry either in the alignment between the poles and radio pulses, or between the poloidal/toroidal structure and the $H\alpha$ emission. In the former case, the simplest possibility is that the field is quadrupolar, leading naturally to the possibility of a $1/4$ phase lag (as opposed to 0 or $1/2$ phase lag for the dipole field). However, the quadrupole cannot produce uniform $H\alpha$ emission since this would result in a non-varying light curve. Instead, two of the quadrupole “lobes” have to produce enhanced $H\alpha$ emission – for example, one each in the northern and southern hemispheres. This configuration is shown in Figure 5, and it appears to explain all of the available observations.

The alternative hypothesis of symmetry breaking can accommodate several possibilities. In the simplest scenario, the radio emission does not emerge uniformly from the polar caps, but is instead produced in a localized region at lower latitude (a “hot spot configuration”). As the object rotates the hot spot moves in and out of our line of sight. This may work in both the toroidal and poloidal dipole configurations, since both can explain the sinusoidal $H\alpha$ variability. However, the arbitrary location of the hot spot does not trivially explain the exact $1/4$ phase lag observed here.

More complex possibilities exist for both the high-order multipole scenario and the broken symmetry scenario. However, the stability of the radio pulses and the overall smooth $H\alpha$ light curve suggest that the dominant topology is not sig-

nificantly more complex than the scenarios we explored here.

To summarize, we find that the $1/4$ phase lag between the radio and $H\alpha$ light curves eliminates simple and symmetric dipole field configurations in which the radio emission is produced at the poles, and the $H\alpha$ emission arises from plasma confined by the dipole field; in this configuration we expect the phase lag to be either zero (toroidal field) or $1/2$ (poloidal field). The simplest alternative is to either: (i) retain the polar origin of the radio emission, but appeal to a quadrupole field with non-uniform $H\alpha$ emission; or (ii) retain the dipole configuration as the origin of the $H\alpha$ emission, but shift the location of the radio emission from the poles to a hot spot at lower latitude. We finally note that in any of the possible configurations, the quiescent non-variable radio emission is likely produced by the largest scale field structure, such that the projection effects caused by the rotation and magnetic axis tilt are minimized.

The possible field topologies that give rise to the observed radio and $H\alpha$ emission appear to be somewhat more complex than those inferred for mid-M dwarfs from spectropolarimetric observations (Donati et al. 2006; Morin et al. 2008). Zeeman-doppler imaging points to predominantly dipole poloidal fields, particularly for objects with the lowest Rossby numbers, $Ro \equiv P/\tau_c \lesssim 0.05$, (τ_c is the convective overturn timescale). For 2M0746+20 we estimate $\tau_c \approx (MR^2/L)^{1/3} \approx 200$ d (§5), so the inferred Rossby number is extremely small, $Ro \approx 4 \times 10^{-4}$, and we may have expected the dipole poloidal configuration to dominate. Thus, our observations indicate that the field topology possibly evolves from mid-M dwarfs to L dwarfs, and that Ro may not be the only relevant parameter.

5. THE MASS-RADIUS RELATION

We next use the rotation period of 2M0746+20 as determined from the radio and $H\alpha$ emission to measure its radius. We make the reasonable assumption that the inclination of the rotation axis, i , is identical to the orbital orientation of $142 \pm 3^\circ$ (Hale 1994). Thus, the projected $v \sin i = 27 \pm 3$ km s $^{-1}$ translates to an actual rotation velocity of $v = 46 \pm 6$ km s $^{-1}$. For a period of 124.32 ± 0.11 min, the corresponding radius of 2M0746+20 is $R = (5.4 \pm 0.7) \times 10^9$ cm, or $0.078 \pm 0.010 R_\odot$. As far as we know, this is the first radius estimate (without a $\sin i$ degeneracy) for an L dwarf through photometric variability.

The total mass of the binary, inferred by Bouy et al. (2004) from the orbital dynamics, is $0.146^{+0.016}_{-0.006} M_\odot$. These authors further estimated a range of $0.075 - 0.095 M_\odot$ for the primary by comparing the infrared photometry with DUSTY model isochrones. Gizis & Reid (2006), on the other hand, argued that the primary mass is $0.078 - 0.082 M_\odot$, primarily as a result of an older age for the system. Taking the conservative approach, we adopt here the wider range of values from Bouy et al. (2004).

The inferred radius and mass are plotted in Figure 6. Also shown are the mass-radius relations from the evolutionary models of Baraffe et al. (1998) for ages of 0.5, 1, and 2 Gyr and solar metallicity, as well as masses and radii for stars in the range $0.1 - 0.7 M_\odot$ from a recent compilation by López-Morales (2007). 2M0746+20a lies below the model predictions, $R \approx 0.105 - 0.115 R_\odot$, by about 30%. This is surprising since stars with known radii in the mass range $0.1 - 0.3 M_\odot$ appear to generally agree with the model predictions (López-Morales 2007). Moreover, stars in the mass range $0.4 - 0.7 M_\odot$ generally lie above the model predic-

tions, possibly as a result of magnetic activity (López-Morales 2007). It is somewhat surprising then, that a highly active L dwarf would underlie the same theoretical models. Similarly, a recent estimate of $R \gtrsim 0.117 R_{\odot}$ for the M8.5 dwarf LSR 1835+32 (Hallinan et al. 2008), based on similar periodic radio emission to that of 2M 0746+20, indicates that there is possibly a larger than expected dispersion in the radii of ultracool dwarfs.

With only a single object it is impossible to assess whether the apparent discrepancy indicates that the theoretical models break down across the sub-stellar boundary, or if this is the result of our assumption that the rotation axis is aligned with the orbital inclination. It is unlikely that the problem is with the value of $\sin i$ since several groups have measured consistent values (§2). If we relax our assumption of $i = 142^{\circ}$, then the minimum allowed radius (corresponding to $i = 90^{\circ}$) is $0.046 R_{\odot}$. The radius corresponding to the 90% probability of the inclination axis distribution ($i > 26^{\circ}$) is $R \approx 0.105 R_{\odot}$ (i.e., there is only a 10% probability that $R > 0.105 R_{\odot}$). This latter value is marginally consistent with the theoretical predictions, but it requires the rotation axis to be severely misaligned with the orbital inclination.

6. SUMMARY AND CONCLUSIONS

We presented simultaneous radio, X-ray, optical, and UV observations of the L dwarf binary 2M 0746+20. Both the radio and $H\alpha$ light curves are dominated by periodic emission, with $P \approx 124$ min. The agreement between the two periods indicates that both arise from the same binary member, which we assume to be the primary star. We note that high angular resolution Very Long Baseline Array observations may pinpoint the origin of the radio emission.

The observed period is well-matched to the rotation velocity of 2M 0746+20, and the assumption that the rotation axis is aligned with the orbital inclination allows us to infer a radius of $R = 0.078 \pm 0.010 R_{\odot}$. Combined with the estimated mass of $0.085 \pm 0.010 M_{\odot}$ (Bouy et al. 2004), we find that the radius is about 30% smaller than predicted from theoretical evolutionary models. It is unclear from this single object whether this is the result of a breakdown in the models, or an indication that the rotation and orbital axes are severely misaligned. Both possibilities have important implications for our understanding of ultracool dwarfs and their formation mechanism.

The combination of periodic radio and $H\alpha$ emission further allows us to explore the magnetic field topology. The radio periodicity is in the form of short duration pulses with a duty cycle of 0.8% and a uniform sense of circular polarization, while the $H\alpha$ emission is sinusoidal and leads the radio emission by 1/4 of a phase. These emission properties can be explained in the context of a rotating mis-aligned magnetic field, but the observed phase lag rules out a symmetric dipole field. Instead, we conclude that either the field is dominated by a quadrupole configuration, or the emission regions are not trivially aligned with a dipole field. In both scenarios the quiescent and non-variable radio emission most likely

arises from the largest scales of the field, which are least affected by rotation and inclination effects.

We note that the availability of simultaneous radio and optical data is crucial to our understanding of the field geometry. If only one band was available to us (as in all previous cases of detected radio pulses), we would have concluded that the dipole topology is the most likely scenario. Instead, the field configuration of 2M 0746+20 appears to be somewhat more complex than those inferred for mid-M dwarfs from phase-resolved optical spectropolarimetry (Donati et al. 2006; Morin et al. 2008), although the vast difference in techniques may result in sensitivity to different field structures. Regardless of the exact configuration, our observations support the general conclusions of recent numerical dynamo simulations (Browning 2008), which predict large-scale fields at low Rossby numbers. Indeed, we estimate that the Rossby number for 2M 0746+20 is very low, $\sim 4 \times 10^{-4}$.

Future observations of 2M 0746+20 will allow us to determine the stability timescale of the magnetic field configuration, a crucial constraint on dynamo models. Previous observations of the L3.5 dwarf 2M 0036+18 (Berger et al. 2005; Hallinan et al. 2008) indicated stability on a timescale of at least ~ 3 yr. In addition, continued surveys for radio emission from low mass stars and brown dwarfs, particularly in binary systems, are warranted. As demonstrated here, these observations, particularly in conjunction with optical spectroscopy, can provide important constraints not only on the mass-radius relation below $\sim 0.1 M_{\odot}$, but also on the magnetic field topology of fully convective objects.

We thank the Chandra, Gemini, VLA, and Swift schedulers for their invaluable help in coordinating these observations. This work has made use of the SIMBAD database, operated at CDS, Strasbourg, France. It is based in part on observations obtained at the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the National Science Foundation (United States), the Science and Technology Facilities Council (United Kingdom), the National Research Council (Canada), CONICYT (Chile), the Australian Research Council (Australia), CNPq (Brazil) and CONICET (Argentina). Data from the UVOT instrument on Swift were used in this work. Swift is an international observatory developed and operated in the US, UK and Italy, and managed by NASA Goddard Space Flight Center with operations center at Penn State University. Support for this work was provided by the National Aeronautics and Space Administration through Chandra Award Number G08-9013A issued by the Chandra X-ray Observatory Center, which is operated by the Smithsonian Astrophysical Observatory for and on behalf of the National Aeronautics Space Administration under contract NAS8-03060.

REFERENCES

- Antonova, A., Doyle, J. G., Hallinan, G., Bourke, S., & Golden, A. 2008, *A&A*, 487, 317
- Antonova, A., Doyle, J. G., Hallinan, G., Golden, A., & Koen, C. 2007, *A&A*, 472, 257
- Audard, M., Osten, R. A., Brown, A., Briggs, K. R., Guedel, M., Hodges-Kluck, E., & Gizis, J. E. 2007, *ArXiv e-prints*, 707
- Bailer-Jones, C. A. L. 2004, *A&A*, 419, 703
- Baraffe, I., Chabrier, G., Allard, F., & Hauschildt, P. H. 1998, *A&A*, 337, 403
- Berger, E. 2002, *ApJ*, 572, 503
- Berger, E. 2006, *ApJ*, 648, 629
- Berger, E., et al. 2001, *Nature*, 410, 338
- Berger, E., et al. 2008a, *ApJ*, 673, 1080
- Berger, E., et al. 2008b, *ApJ*, 676, 1307
- Berger, E., et al. 2005, *ApJ*, 627, 960
- Bouy, H., et al. 2004, *A&A*, 423, 341
- Browning, M. K. 2008, *ApJ*, 676, 1262

- Burgasser, A. J., & Putman, M. E. 2005, *ApJ*, 626, 486
- Dahn, C. C., et al. 2002, *AJ*, 124, 1170
- Donati, J.-F., Forveille, T., Cameron, A. C., Barnes, J. R., Delfosse, X., Jardine, M. M., & Valenti, J. A. 2006, *Science*, 311, 633
- Gizis, J. E., & Reid, I. N. 2006, *AJ*, 131, 638
- Güdel, M. 2002, *ARA&A*, 40, 217
- Hale, A. 1994, *AJ*, 107, 306
- Hallinan, G., Antonova, A., Doyle, J. G., Bourke, S., Briskin, W. F., & Golden, A. 2006, *ApJ*, 653, 690
- Hallinan, G., Antonova, A., Doyle, J. G., Bourke, S., Lane, C., & Golden, A. 2008, *ArXiv e-prints*, 805
- Hallinan, G., et al. 2007, *ApJ*, 663, L25
- Hook, I. M., Jørgensen, I., Allington-Smith, J. R., Davies, R. L., Metcalfe, N., Murowinski, R. G., & Crampton, D. 2004, *PASP*, 116, 425
- López-Morales, M. 2007, *ApJ*, 660, 732
- Morin, J., et al. 2008, *ArXiv e-prints*, 808
- Osten, R. A., Hawley, S. L., Bastian, T. S., & Reid, I. N. 2006, *ApJ*, 637, 518
- Phan-Bao, N., Osten, R. A., Lim, J., Martín, E. L., & Ho, P. T. P. 2007, *ApJ*, 658, 553
- Reid, I. N., Kirkpatrick, J. D., Gizis, J. E., Dahn, C. C., Monet, D. G., Williams, R. J., Liebert, J., & Burgasser, A. J. 2000, *AJ*, 119, 369
- Reid, I. N., Kirkpatrick, J. D., Liebert, J., Gizis, J. E., Dahn, C. C., & Monet, D. G. 2002, *AJ*, 124, 519
- Reiners, A., & Basri, G. 2006, *ApJ*, 644, 497
- Reiners, A., & Basri, G. 2007, *ApJ*, 656, 1121
- Reiners, A., & Basri, G. 2008, *ArXiv e-prints*, 805
- Schmidt, S. J., Cruz, K. L., Bongiorno, B. J., Liebert, J., & Reid, I. N. 2007, *AJ*, 133, 2258
- Treumann, R. A. 2006, *A&A Rev.*, 13, 229
- Vrba, F. J., et al. 2004, *AJ*, 127, 2948

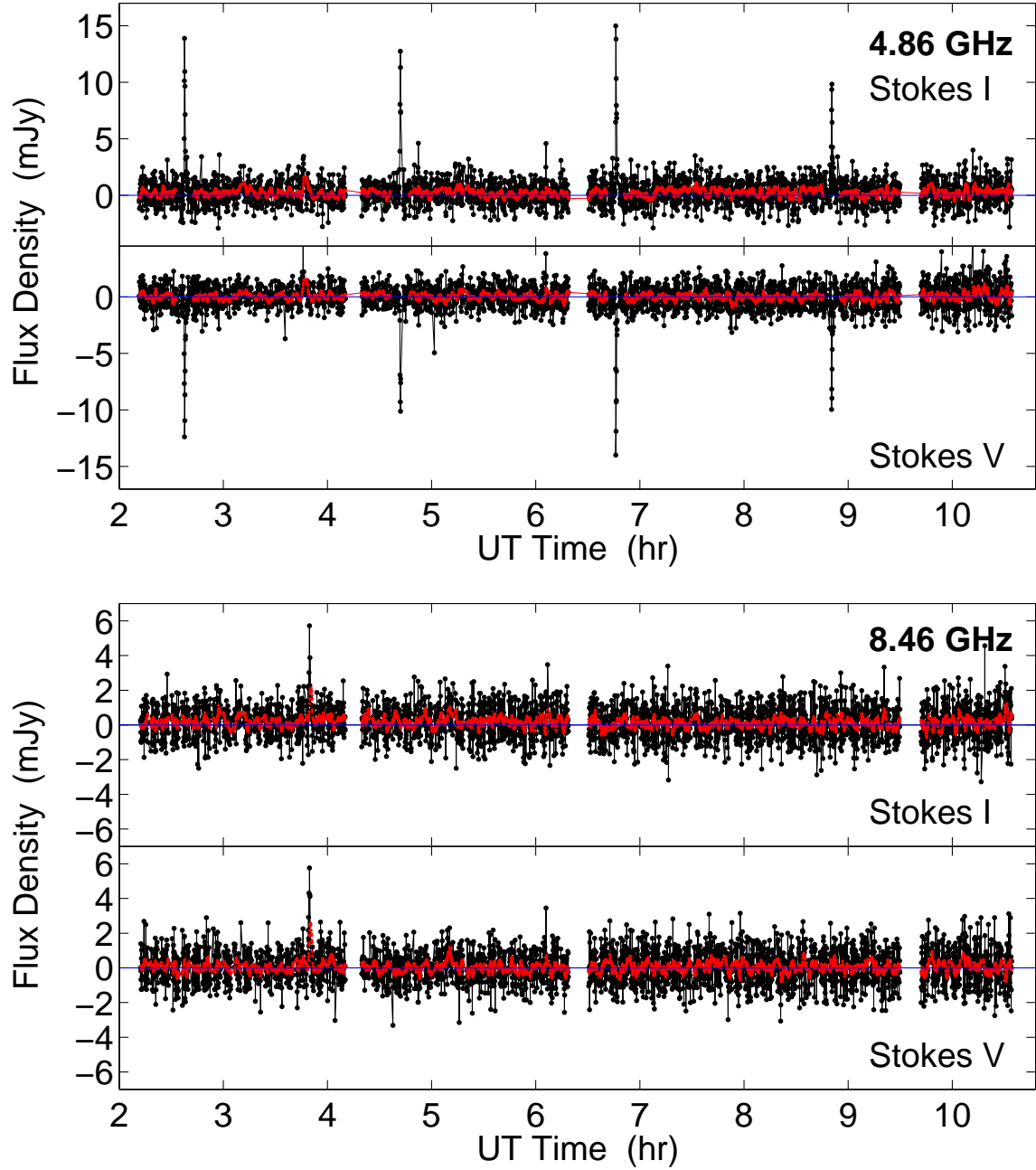


FIG. 1.— Radio light curves of 2M 0746+20 at 4.86 GHz (top) and 8.46 GHz (bottom) in stokes I (total intensity) and stokes V (circular polarization; negative values indicate left-handed polarization). The data are shown at the native 5-s time resolution (black points) and smoothed with a 1-min boxcar (red points). The radio emission at 4.86 GHz is dominated by bright, $\sim 100\%$ circularly polarized, and periodic bursts ($P = 124.32 \pm 0.11$ min), with no corresponding emission at 8.46 GHz. These properties point to coherent emission. In addition, we detect quiescent, non-variable emission at both frequencies, with a spectral index indicative of gyrosynchrotron radiation.

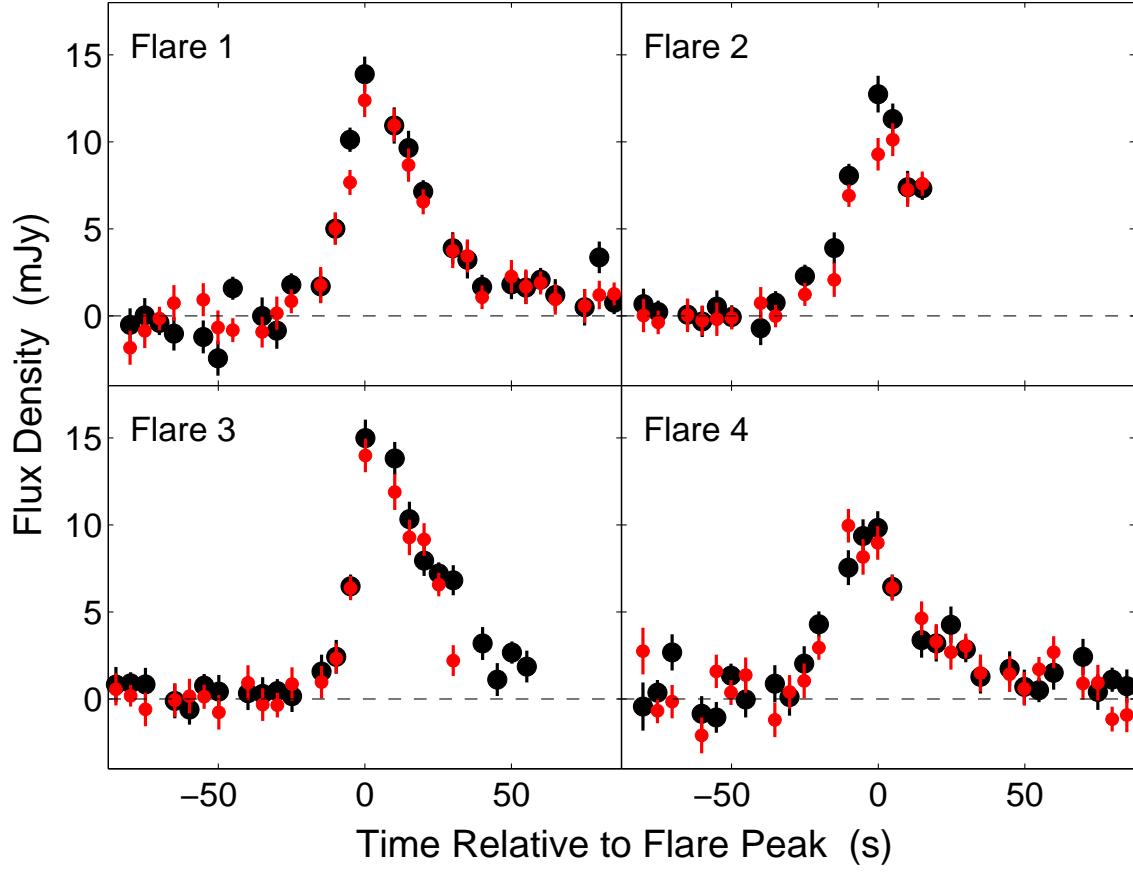


FIG. 2.— Zoom-in view on the four bursts detected at 4.86 GHz. Time is measured relative to the peak of each burst. Both total intensity (black) and circular polarization (red) are shown, with the latter inverted for ease of comparison. The bursts are nearly 100% circularly polarized. The rise time of the bursts is about 20 s, while the decay time is about 40 s.

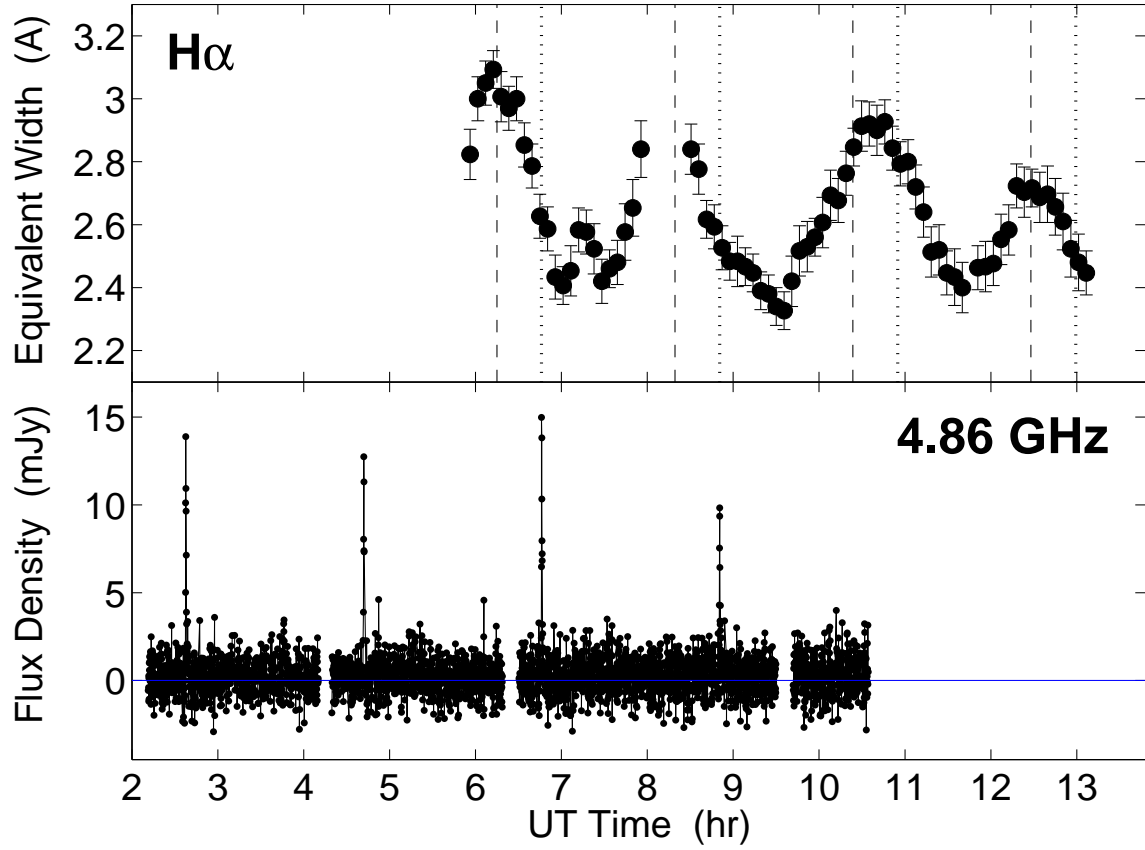


FIG. 3.— $H\alpha$ equivalent width light curve (top) in comparison to the 4.86 GHz total intensity light curve (bottom). The $H\alpha$ emission is clearly sinusoidal, with the same period as that measured from the radio bursts. The peaks of the $H\alpha$ light curve (dashed lines) lead the radio bursts (dotted lines) by exactly 1/4 of a period. These properties, along with the non-zero minima, point to emission from a large-scale chromospheric structure whose projected solid angle varies with the rotation of 2M 0746+20.

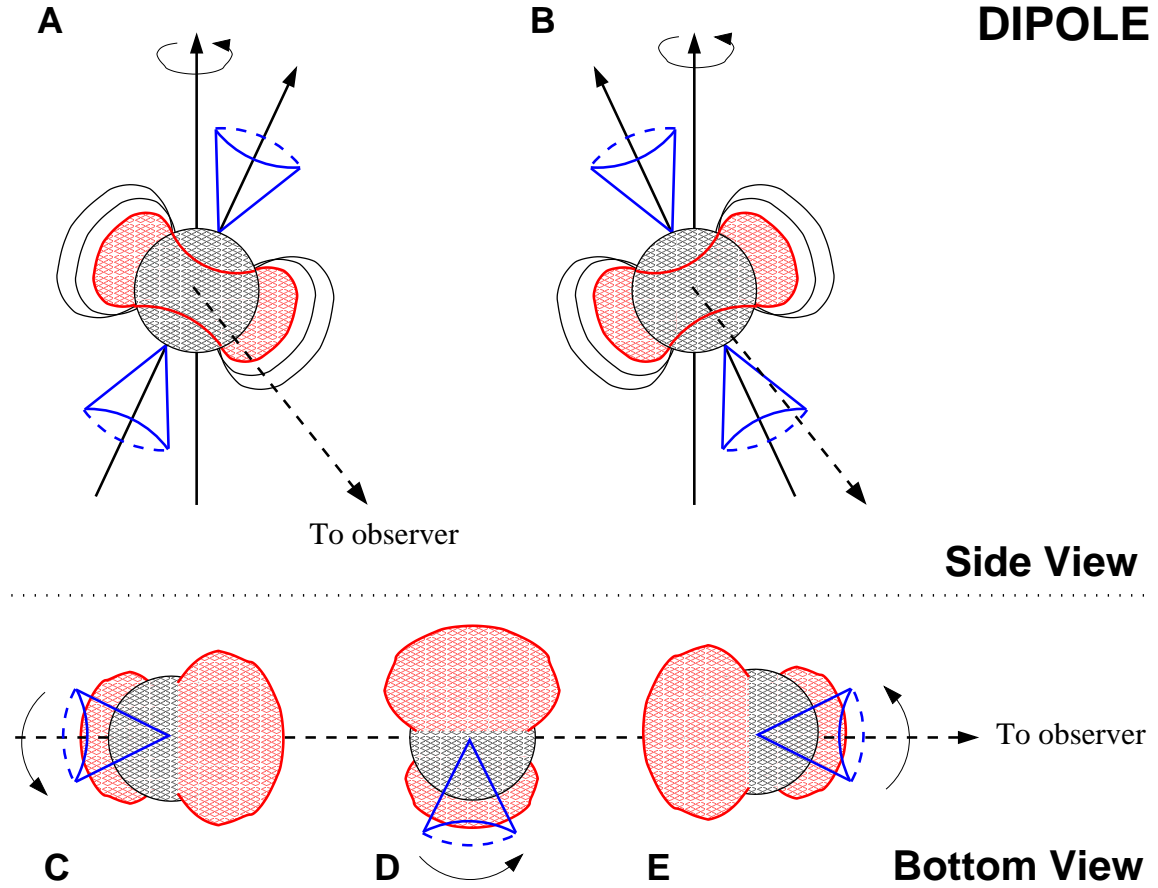


FIG. 4.— A dipole poloidal field topology for 2M 0746+20, and its orientation relative to the line of sight. This configuration leads to radio bursts (blue) and sinusoidal $H\alpha$ emission (red). However, as shown in both sets of projections, we expect a $1/2$ phase lag between the two bands, with the minimum projected solid angle of the field (i.e., minimum $H\alpha$ emission) coinciding with the radio bursts. In the alternative toroidal configuration the $H\alpha$ and radio peaks will be aligned since both are produced in the same polar region.

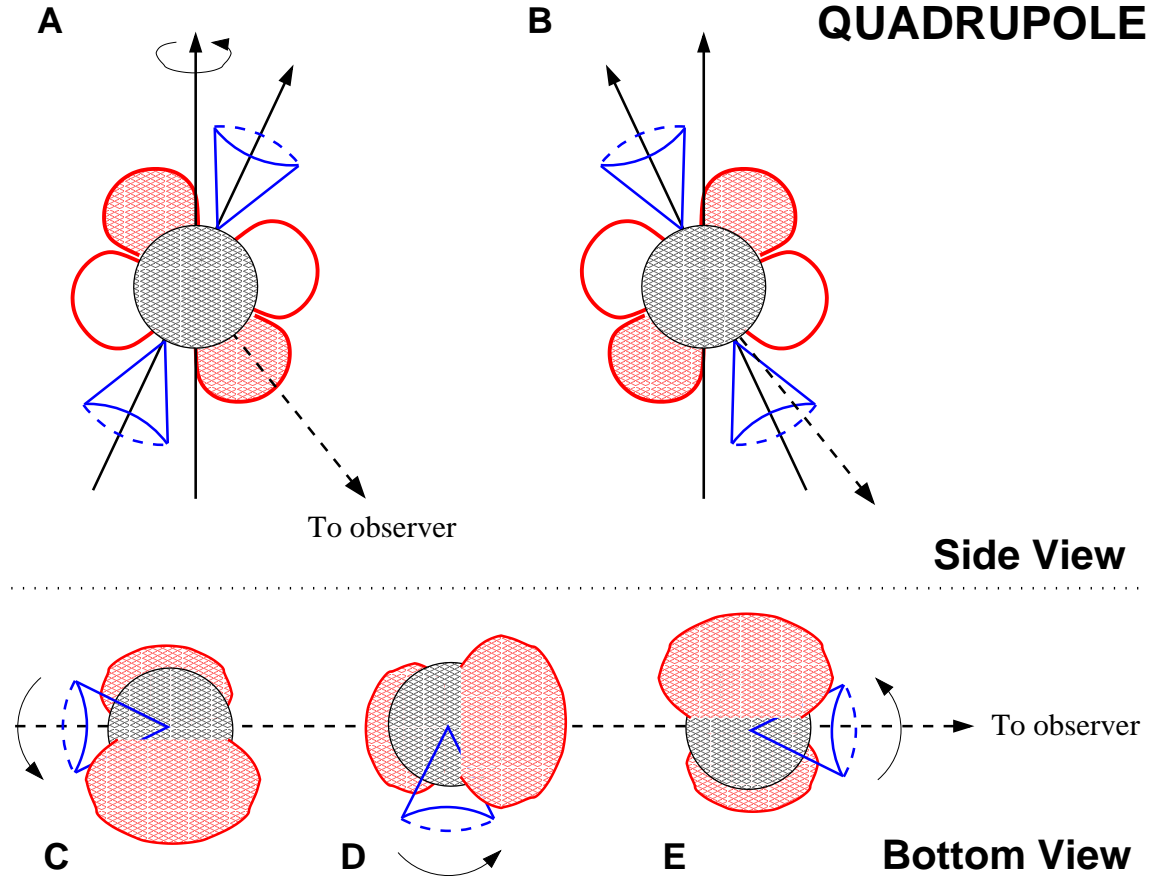


FIG. 5.— A quadrupole field configuration for 2M0746+20, and its orientation relative to the line of sight. This configuration reproduces both the light curve shapes and the $1/4$ phase lag. However, it requires non-uniform $H\alpha$ surface brightness, with two of the quadrupole lobes (on opposite hemispheres) producing stronger $H\alpha$ emission.

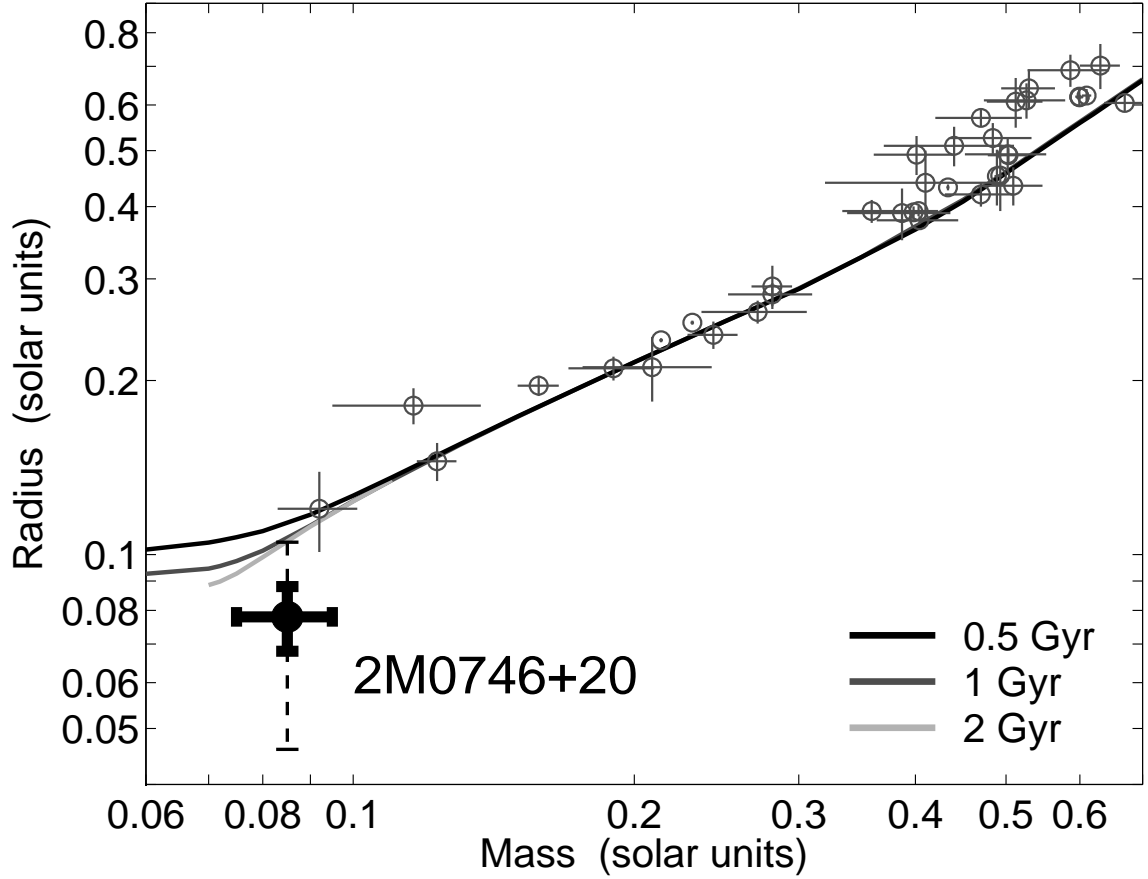


FIG. 6.— Stellar radius as a function of mass for the primary star in 2M 0746+20. Also shown are model tracks (Baraffe et al. 1998) for stellar ages of 0.5, 1, and 2 Gyr, as well as a compilation of higher mass stars (open symbols; López-Morales 2007). The inferred radius of 2M 0746+20a is about 30% smaller than predicted by the models, with the smallest discrepancy for the oldest age. The overall disagreement indicates either a problem with the evolutionary models near the bottom of the main sequence, or that our assumption that the rotation and orbital inclinations are identical is wrong. The dashed line indicates the range of inferred radii that corresponds to the minimum allowed radius (i.e., $i = 90^\circ$) and the 90% inclination probability (i.e., there is $< 10\%$ probability that the inclination angle is such that the inferred radius will be larger than this value). This latter value is marginally consistent with the theoretical models.